

Manifold invariants affect dynamics in ADS gravity

Tomáš Liko*

Department of Mathematical and Statistical Sciences

University of Alberta

Edmonton, AB T6G 2G1, Canada

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Abstract

The first-order Holst action with negative cosmological constant is rendered finite by requiring functional differentiability on the configuration space of tetrads and connections. The surface terms that arise in the action for ADS gravity are equivalent to the Euler and Pontryagin densities with fixed weight factors; these terms modify the Noether charges that arise from diffeomorphism invariance of the action.

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Let \mathcal{M} be a four-dimensional manifold with boundary $\partial\mathcal{M}$. The configuration space \mathcal{C} of the gravitational field is the pair $\{e, A\}$, consisting of the coframe $e^I := e_a^I dx^a$ and spin connection $A^{IJ} := A_a^{IJ} dx^a$ (with spacetime indices $a, b, \dots \in \{0, \dots, 3\}$ and internal indices $I, J, \dots \in \{0, \dots, 3\}$). In terms of e and A , the spacetime metric and curvature 2-form are (resp.) $g_{ab} = \eta_{IJ} e_a^I \otimes e_b^J$ and $\Omega^I_J = dA^I_J + A^I_K \wedge A^K_J = (1/2)R^I_{JKL} e^K \wedge e^L$, with R^I_{JKL} being the Riemann tensor.

The first-order Holst action with negative cosmological constant $\Lambda := -3\ell^{-2} \in \mathbb{R}_-$ and Barbero-Immirzi parameter $\gamma \in \mathbb{R} \setminus \{0\}$ on a four-dimensional manifold (\mathcal{M}, e, A) with generic boundary $\partial\mathcal{M}$ is [1–4]:

$$I = \frac{1}{32\pi} \int_{\mathcal{M}} \epsilon_{IJKL} \left(e^I \wedge e^J \wedge \Omega^{KL} + \frac{1}{2\ell^2} e^I \wedge e^J \wedge e^K \wedge e^L + \frac{\ell^2}{2} \Omega^{IJ} \wedge \Omega^{KL} \right) - \frac{1}{\gamma} (e_I \wedge e_J \wedge \Omega^{IJ} + \ell^2 \Omega_{IJ} \wedge \Omega^{IJ}) . \quad (1)$$

This form of the action can be used to derive the Noether charges associated with diffeomorphisms. Specifically, for a smooth vector field ξ generating a spacetime diffeomorphism, the symplectic current 3-form is $\star\mathcal{J} = \Theta(e, A, \delta A) - I_\xi \mathcal{L}$, with Θ being the (first) variation of the surface terms in (1), \mathcal{L} the Lagrangian density in

*Electronic mail: tliko@math.ualberta.ca

(1), \star denoting Hodge duality and I_ξ denoting the contraction operator acting on p -forms. For the Holst-ADS action, the symplectic current is a total derivative, and can be integrated over a two-dimensional spatial cross section Σ of $\partial\mathcal{M}$. The resulting conserved charge that arises from diffeomorphism invariance of the first-order Holst-ADS action is:

$$\mathcal{Q}[\xi] = \frac{\ell^2}{32\pi} \oint_\Sigma \epsilon_{IJKL} \bar{\Omega}^{IJ} \wedge I_\xi A^{KL} - \frac{1}{\gamma} \bar{\Omega}_{IJ} \wedge I_\xi A^{IJ}, \quad (2)$$

with $\bar{\Omega} := \Omega + (1/\ell^2)e \wedge e$. This charge is an extension of the Noether charge found by Aros *et al* [1] to include a real and non-zero Barbero-Immirzi parameter. The presence of γ in the action therefore changes the Noether charge for ADS spacetimes. This contribution to the charge vanishes for spacetimes that have diagonalizable metrics or zero Pontryagin number [2–4].

Example 1: Euler density and black-hole entropy. In this example, we consider spacetimes containing a (local) Killing horizon Δ as inner boundary and an asymptotic region \mathcal{I} as outer boundary. The topology of Δ is $\mathbb{R} \times \Sigma_g$, with Σ_g being a compact 2-manifold with genus $g \in \mathbb{Z}_+ \cup \{0\}$. The class of spacetimes thus include the Schwarzschild-ADS and Kerr-ADS families of solutions describing black holes with ‘topological’ horizons (see, e.g. Galloway *et al* [5] and references therein).

For the class of spacetimes under consideration, the Euler term is non-zero and the Pontryagin term is zero in the *on-shell* first-order action. A non-zero Euler term in the action modifies the entropy of the corresponding black holes. The entropy for these spacetimes that satisfies the first law of black-hole mechanics is $\mathcal{S} := \mathcal{A}/4 + \pi\ell^2(1 - g)$, with \mathcal{A} being the surface area of the two-dimensional cross section Σ_g of the horizon [6, 7]. The toroidal horizon is a special case, for which the entropy is $\mathcal{S} = \mathcal{A}/4$.

Let us consider the merging of two black holes. The entropies of these black holes are $\mathcal{S}_1 = (1/4)[\mathcal{A}_1 + 4\pi\ell^2(1 - g_1)]$ and $\mathcal{S}_2 = (1/4)[\mathcal{A}_2 + 4\pi\ell^2(1 - g_2)]$. If the total entropy before the black holes merge is $\mathcal{S} = \mathcal{S}_1 + \mathcal{S}_2 = (1/4)[\mathcal{A}_1 + \mathcal{A}_2 + 4\pi\ell^2(2 - g_1 - g_2)]$, and the total entropy of the final black hole configuration is $\mathcal{S}' = (1/4)[\mathcal{A}' + 4\pi\ell^2(1 - g')]$, the second law of thermodynamics will hold if and only if $\mathcal{S}' > \mathcal{S}_1 + \mathcal{S}_2$. Thus we have the following upper bound on the ADS radius ℓ :

$$\ell < \left[\left(\frac{1}{4\pi} \right) \left(\frac{\mathcal{A}' - \mathcal{A}_1 - \mathcal{A}_2}{1 + g' - g_1 - g_2} \right) \right]^{1/2}. \quad (3)$$

The reality of ℓ implies that $\mathcal{A}' > \mathcal{A}_1 + \mathcal{A}_2$ *if and only if* $g' > g_1 + g_2 - 1$. In other words, the number of genera cannot decrease in a physical process for which the second law of thermodynamics (in the form $\mathcal{S}' > \mathcal{S}_1 + \mathcal{S}_2$) *and* the second law of black-hole mechanics (in the form $\mathcal{A}' > \mathcal{A}_1 + \mathcal{A}_2$) hold together. Thus, for example, two toroidal black holes cannot merge into a final toroidal configuration because the number of genera of the final state has to be at least 2. ■

Remarks. (i) The bound (3) is consistent with area bounds that were found by Gibbons [8] and by Woolgar [9]. For black holes in ADS spacetime, the surface area of the black hole is constrained by the bound: $\mathcal{A} \geq 4\pi(g-1)/|\Lambda|$. This bound says that higher-genus horizons necessarily have greater areas. (ii) Similar area bounds can be obtained for black holes in spacetimes with *positive* cosmological constant. For black holes merging in de Sitter spacetimes, it turns out that the second law of thermodynamics holds while the second law of black-hole mechanics is *violated*.

Example 2: Pontryagin density and gravitational entropy. In this example, we consider a more general class of spacetimes than the ones considered in Example 1. For these spacetimes, the Pontryagin number is taken to be non-zero. Spacetimes with non-zero Pontryagin number include the Taub-NUT solution [10]. The entropy of the Taub-NUT-ADS spacetime is found to be [3, 4]

$$\mathcal{S} = \mathcal{S}_0 + \frac{4\pi N^2}{\gamma} \left(1 - \frac{N^2}{\ell^2}\right). \quad (4)$$

Here, \mathcal{S}_0 is the entropy of the Taub-NUT-ADS solution without the Holst term [1, 11, 12]. ■

Summary. Topological invariants of a manifold with boundary, although total derivatives in the action, affect the dynamics of black objects in four dimensions.

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